Extremely Long-Lived Charged Massive Particles as A Probe for Reheating of the Universe

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WMAP DATA (First year)

$\Omega_{DM} = 0.23 \pm 0.04$

$\Omega_{Total} = 1.02 \pm 0.02$

$\Omega_{baryon} = 0.044 \pm 0.004$

More than 90 % of total energy of the universe is unknown

Dark matter ～ 20 %
Dark energy ～ 70 %

What is the dark matter ? Why is the relic $\Omega \sim O(0.1)$?
How were they generated ? When ?…..

Extremely Long-Lived Charged Massive Particle

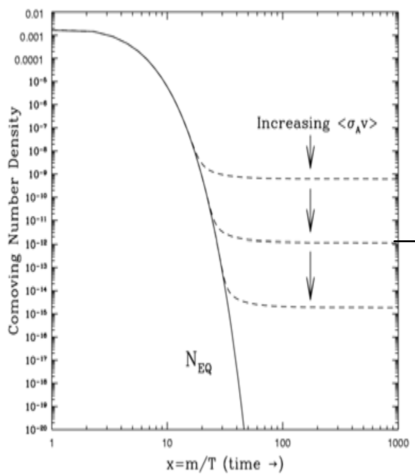
TeV scale new physics (SUSY, Extra Dimensions)

→ Copies of SM particles → Lightest one ~ stable

Parity SM(even) Copies(odd)

Lightest one in SM sector : charged slepton, charged KK leptons etc

SuperWIMP dark matter scenario motivates Extremely Long-Lived CHAMPs.



J.Feng, A.Rajaraman, F.T(2003), J.Feng,S.Su,F.T(2004)

e.g Lightest one : Gravitino LSP,

Lightest one in SM : Stau NLSP(Selectron)

(Spanos talk)

Decay after BBN

NR decay: $m_{NLP} \sim m_{superWIMP}$

$\Omega_{superWIMP} \sim (m_{superWIMP}/m_{CHAMP})\Omega_{CHAMP}$

SuperWIMP dark matter → mass degeneracy

。。。Late decay may not always provide the leading effects in the early Universe

Other phenomenon in early universe might happen without significant change from SBBN due to late decays.

Prospects of collider experiments for extremely long lived CHAMP search

Discovery (Heavily Ionizing Track, TOF etc)

M.Drees, X.Tata(1990),J.Goity,W.Kossler,M.Sher(1993)J.Feng,T.Moroi(1998)

Tevatron $m_C \sim 180$GeV ($L=10fb^{-1}$, stable stau inside collider detector)

Mass, Couplings with SM particles, Lifetime, Decay properties

Trapping CHAMPs

B.T.Smith, J.Feng(2004) K.Hamaguchi,Y.Kuno,T.Nakaya,M.Nojiri(2004)

CBBN

K.Kohri, F.T (2006), M.Pospelov (2006), M.Kaplinghat, A.Rajaraman(2006),
R.Cyburt,J.Ellis,B.Field,K.Olive,V.Spanos(2006)

Bound state of a light element and a CHAMP during/after BBN

Kohri, F.T(2006)

Bound state formation → Photo destruction

$$[\frac{\partial}{\partial t}n_X]_{\rm capture} \simeq - < \sigma_r v > (n_C n_X - n_{(C,X)} n_\gamma (E > E_{\rm bin}))$$

$$n_\gamma (E > E_{\rm bin}) \equiv n_\gamma \frac{\pi^2}{2\zeta(3)} (\frac{m_X}{2\pi T})^{3/2} e^{-\frac{E_{\rm bin}}{T}}$$

$$n_\gamma = \frac{2\zeta(3)}{\pi^2} T^3$$

$$\longrightarrow \quad T_c \simeq \frac{E_{\rm bin}}{40}$$

Nucleus(X)	binding energy (MeV)	atomic number
H	0.025	Z=1
D	0.050	Z=1
T	0.075	Z=1
^{3}He	0.270	Z=2
^{4}He	0.311	Z=2
^{5}He	0.431	Z=2
^{5}Li	0.842	Z=3
^{6}Li	0.914	Z=3
^{7}Li	0.952	Z=3
^{7}Be	1.490	Z=4
^{8}Be	1.550	Z=4
^{10}B	2.210	Z=5

R.N.Cahn,S.L.Glashow(1981)

Heavier elements may be captured in earlier time.

$T_c(^7\mathrm{Be}) \sim 37$keV, $T_c(^7\mathrm{Li}) \sim 25$ keV

SBBN process completely decouple at T ~ 50-20keV

All exponential suppression is significant at below this T

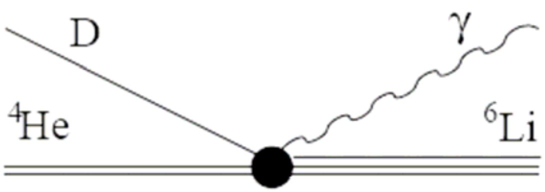
Coulomb suppression (Low E)
Boltzmann suppression (low T)
β decay of neutron etc (small Hubble rate)

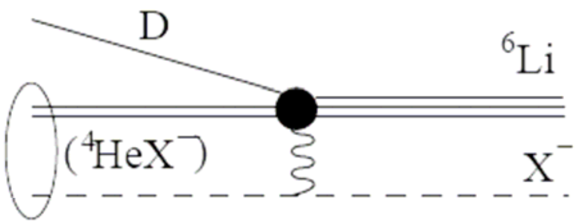
$\tau_n = 885.7 \pm 0.8$ s

The abundance of heavier than Li may be changed from SBBN value.

Virtual Photon processes (M.Pospelov(2006))

$$\mathrm{SBBN}: \ {}^4\mathrm{He} + \mathrm{D} \to {}^6\mathrm{Li} + \gamma; \quad Q = 1.47\mathrm{MeV}$$
$$\mathrm{CBBN}: \ ({}^4\mathrm{He}X^-) + \mathrm{D} \to {}^6\mathrm{Li} + X^-; \ Q \simeq 1.13\mathrm{MeV}$$





→Significant ^{6}Li production relative to the SBBN case

(Lifetime >> 10^4 sec)

: cross section ~ O(10^6) enhancement

E [keV]	$\sigma_{1\to 2}$ [barn]	S [MeV barn]
10	3.85×10^{-6}	0.0426
20	1.09×10^{-4}	0.0410
36.4	6.88×10^{-4}	0.0380
50	1.41×10^{-3}	0.0357
100	3.50×10^{-3}	0.0286

K.Hamaguchi,T. Hatsuda, M.Kamimura, Y.Kino,T.Yanagida(2007)

CBBN and primordial ^{6}Li abundance

$$\frac{dn_{^4\mathrm{He}}}{dt} + 3Hn_{^4\mathrm{He}} = -<\sigma_{\mathrm{rec}} v> (n_C n_{^4\mathrm{He}} - n_{(C,^4\mathrm{He})}\tilde{n}_\gamma) + \frac{1}{\tau_C} n_{(C,^4\mathrm{He})}$$

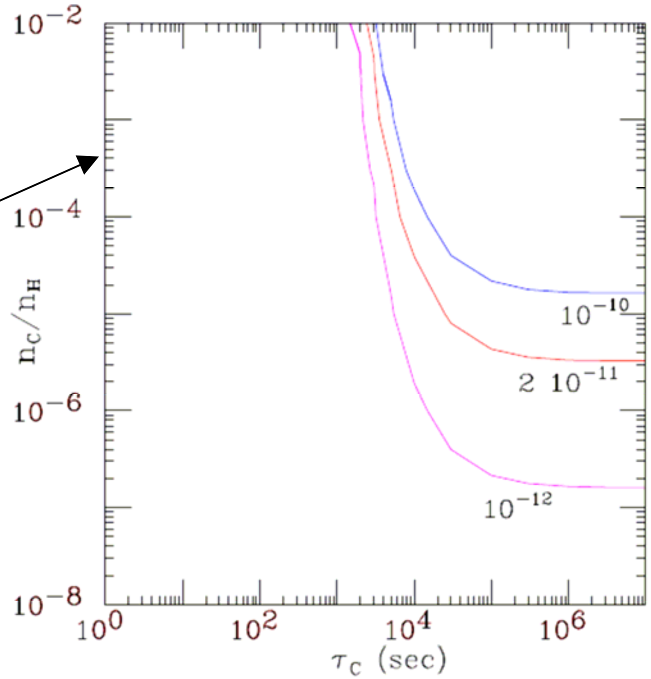
$$\frac{dn_C}{dt} + 3Hn_C = -<\sigma_{\mathrm{rec}} v> (n_C n_{^4\mathrm{He}} - n_{(C,^4\mathrm{He})}\tilde{n}_\gamma) + <\sigma^{CBBN}_{^6\mathrm{Li}} v> n_{(C,^4\mathrm{He})} n_D - \frac{1}{\tau_C} n_C$$

$$\frac{dn_{(C,^4\mathrm{He})}}{dt} + 3Hn_{(C,^4\mathrm{He})} = <\sigma_{\mathrm{rec}} v> (n_C n_{^4\mathrm{He}} - n_{(C,^4\mathrm{He})}\tilde{n}_\gamma) - <\sigma^{CBBN}_{^6\mathrm{Li}} v> n_{(C,^4\mathrm{He})} n_D - \frac{1}{\tau_C} n_{(C,^4\mathrm{He})}$$

$$\frac{dn_{^6\mathrm{Li}}}{dt} + 3Hn_{^6\mathrm{Li}} = <\sigma^{CBBN}_{^6\mathrm{Li}} v> n_{(C,^4\mathrm{He})} n_\mathrm{D}$$

$$\frac{dn_\mathrm{D}}{dt} + 3Hn_\mathrm{D} = -<\sigma^{CBBN}_{^6\mathrm{Li}} v> n_{(C,^4\mathrm{He})} n_\mathrm{D}$$

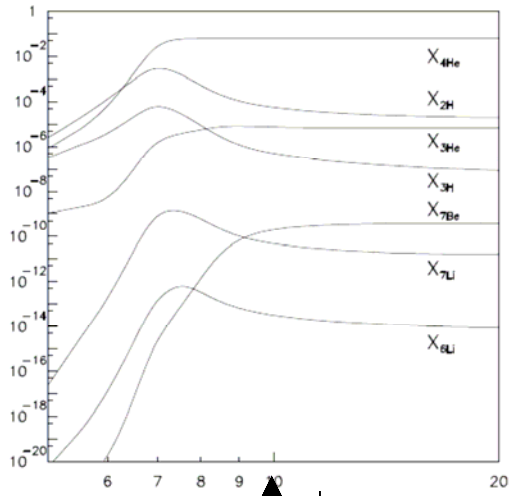
Standard m_C =100GeV

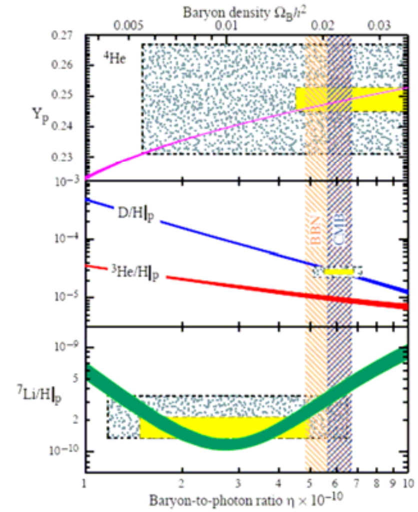


F.T(2007)

The fate of the extremely long-lived CHAMPs

SBBN

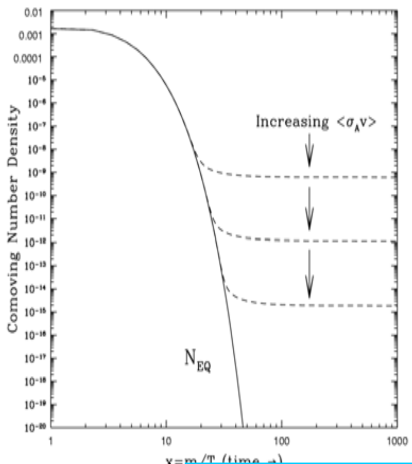




PDG2006

CBBN

CHAMPs chemical decoupling



T ~ 8keV

50keV

CHAMPlate deacy : DM production?

$\Omega_{DM} \sim (m_{DM}/m_{CHAMP})\Omega_{CHAMP}$??

Low reheating temperature??
(Entropy production??)
→hidden parameter
for collider measurements

Matter-Radiation equality

T

$T_f \sim m/25$

1MeV-50keV

T~keV
(Gravitino LSP)

0.1eV

Probe for the Early Universe

Find physical observable which does not significantly change from the primordial values set in the early stage of the history of the Universe.
Or Specify key initial conditions to describe the present universe which can be confirmed by astrophysical observations.
→Light element abundances (He,D,Li…) ^{6}Li

Start from simple cosmological models
and get some inputs from collider experiments.
→Low reheating temperature model

Compare the theoretical prediction with observed values
…Extract hidden parameters in collider experiment alone
→Reheating temperature (Decay rate of key particle for reheating)

Connecting ^{6}Li with the reheating temperature

Low reheating temperature models $(T_{RH} / m_{CHAMP} << O(10))$

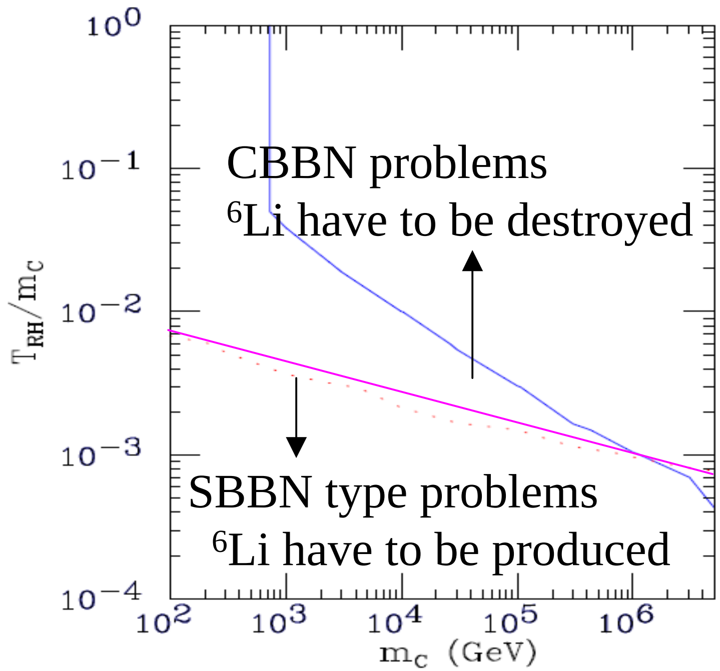
$$\frac{d\rho_\phi}{dt} + 3H\rho_\phi = -\Gamma_\phi \rho_\phi$$
$$\frac{d\rho_R}{dt} + 4H\rho_R = \Gamma_\phi \rho_\phi + <\sigma v> 2 <E_C> [n_C^2 - n_{\rm EQ}^2]$$
$$\frac{dn_C}{dt} + 3Hn_C = - <\sigma v> [n_C^2 - n_{\rm EQ}^2] \qquad (2)$$

Guidice, Kolb, Riotto(2000)

$$\Gamma_\phi = \sqrt{\frac{4\pi^3 g_*(T_{\rm RH})}{45}} \frac{T_{\rm RH}^2}{M_{\rm pl}},$$

$$\Omega_C^{TH(Low)} h^2 = 3.3 \times 10^{-8} [\frac{g_*(T_{\rm RH})^{1/2}}{g_*(T_F)}] \frac{T_{\rm RH}^3 {\rm GeV}^{-2}}{\gamma m_C x_F^{-4}}$$

$\tau_c > 10^4$ sec



F.T(2007)

$\rho_R \sim 0$, $\rho_C \sim 0$ at $t=t_I$

(Other initial condition:Late time Entropy production J.Pradler, F.Steffen(2006))

No direct decay to CHAMPs is assumed. (Inclusion of the decay → Kohri, Yamaguchi, Yokoyama(2005))

Implications of low reheating temperature of the Universe

Difficulties ??

Split hidden sector(Inflaton?) and visible sector (SM sector)

e.g SUGRA (assuming direct reheating from Inflaton)

Gravitino overproduction vs low reheating temperature

Interesting ??

Multi-step energy transfers

Inflaton → Gravitational particles → SM particles

→ Stringly reheating ??

Multi-step reheating

Inflaton→ Radiation Dominated→Matter Dominated→Radiation Dominated

(Late time Entropy production J.Pradler, F.Steffen(2006))

Summary

We discussed extremely long lived charged massive particles as a probe of the early Universe.

Primordial ^{6}Li abundance may be sensitive to the number density of CHAMPs if the lifetime is longer than 10^4 sec.

Discovery of CHAMPs and the measurement of lifetime may provide us some information about the reheating of the Universe for simple cosmological models.

To make my statement robust, we need further efforts to understand the relation between primordial and observed ^{6}Li abundances and to fix uncertainties of CBBN prediction.